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the six- or seven-fold greater average velocity of planetary nebulæ.

A further significant result is the accordance of the value of the solar motion derived from these nebular velocities with that of 20.2 kilometers derived from Class B stars: namely, 20.1 kilometers from the 73 nebulæ of regular form, and 20.7 kilometers from the 7 extended nebulæ. It is thus clear that the extended and planetary nebulæ are associated, at least in respect to their peculiar motions, with our stellar system. The evidence of association of the nebulæ with our stellar system is strengthened by strong indications of preferential motions in the direction of the vertices of KAPTEYN'S two star streams. This is shown in two ways, first by a tabulation with respect to distance from the vertices and second with respect to distance from the central plane of the Milky Way.

W. W. CAMPBELL,
J. H. MOORE.

NOTE ON HIND'S VARIABLE NEBULA.

This nebula (N. G. C. 1554), close to the irregularly variable star T *Tauri,* is an object of special interest in that it affords us probably the only absolutely authentic evidence of change in this class of objects. It was discovered by HIND in 1852, and was seen easily and measured by STRUVE, D'ARREST, LASSELL, and others, but had disappeared by October, 1861. The various observations are so definite that there is no chance for error. D'ARREST, for instance, describes it as quite bright, about 1′ in diameter, and saw it easily in the light of the five-day-old moon. In 1890, Professor BARNARD described it as at the limit of visibility with the 36-inch refractor.

Five plates have been taken of this region with the Crossley Reflector; the dates and exposures are:—

1899, December 6th, exposure	3^h 52^m
1899, December 27th, "	4 00
1914, March 27th, "	1 20
1914, November 16th, "	3 00
1914, November 20th, "	2 30

The nebula is of the utmost faintness on all these plates, but is apparently a little brighter in 1914. It is difficult to resist the impression that it has become somewhat brighter, tho the rather poor character of the plates of 1899 makes a definite decision impossible.

The star T *Tauri* itself has a very interesting nebulous wing which is either variable or must rotate so as to be concealed at times. This appendage is a small cone-shaped projection perhaps four seconds of arc in length in position angle 150 ± °. It was often seen by TEMPEL, and was fairly conspicuous when examined by BARNARD and BURNHAM in 1890, but entirely invisible when BARNARD looked for it in February, 1895. It can be clearly made out in the elongated image of T *Tauri* on both the Crossley plates of 1899. On the 1914 plates it is certainly not present; the disk of the variable is perfectly round, tho perhaps with just a trace of nebulosity all around it. It is evident that this interesting variable will repay further investigation.
H. D. CURTIS.

SEARCH FOR FAINT MEMBERS OF THE TAURUS CLUSTER.

Dr. HERTZSPRUNG suggested in 1914 that it would be of interest to compare the earlier plates of the region of T *Tauri* with recent plates in the endeavor to find very faint members of the well-known Taurus stream. This region is about five degrees from the center of the *Hyades,* and a proper motion of 1″.6 would be expected in the interval of nearly fifteen years. A motion of this amount should show clearly on the Crossley plates.

The measures were made entirely differential, with respect to a common system of five stars of eighth to ninth magnitude; and twelve faint stars of from fourteenth to eighteenth magnitudes were selected at random in the area of best definition. None of the small relative motions found agree either in magnitude or direction with the motion of the components of the Taurus stream. Within the limited area examined the results must be regarded as purely negative.

H. D. CURTIS.

NOTE ON COMET *d* 1915 (MELLISH).

Comet *d* 1915 was first seen by Mr. JOHN E. MELLISH, of the Yerkes Observatory, on September 13, 1915.

The first observation received at the Students' Observatory